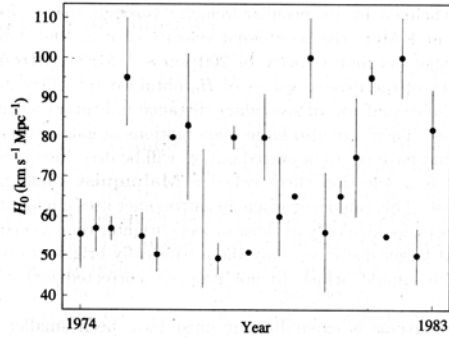
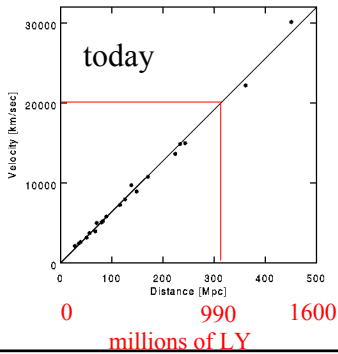
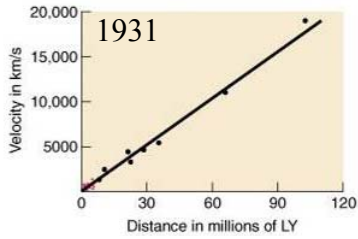


Hubble's Law



Little h

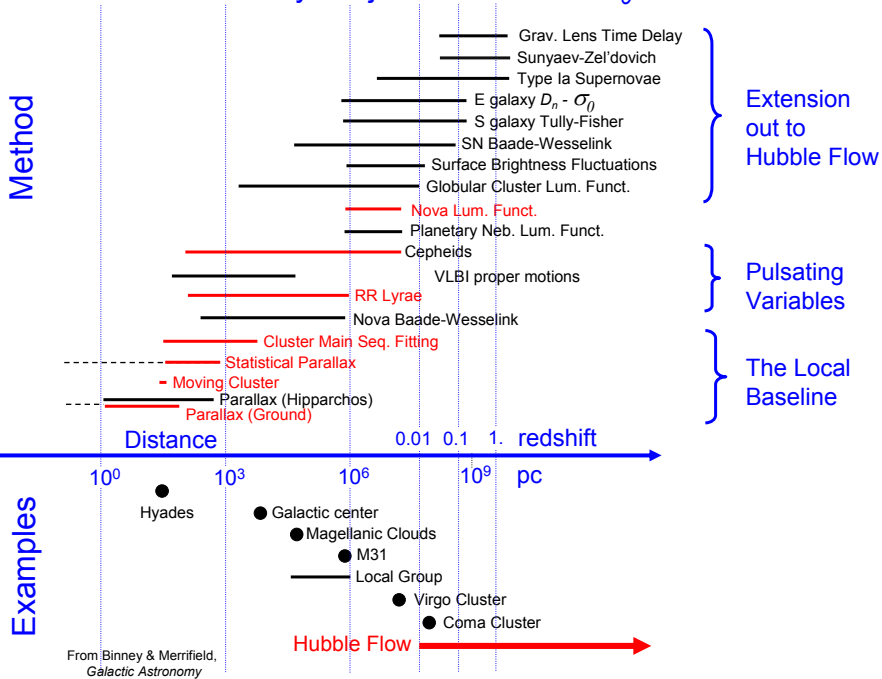
- The Hubble Un-constant (blush)

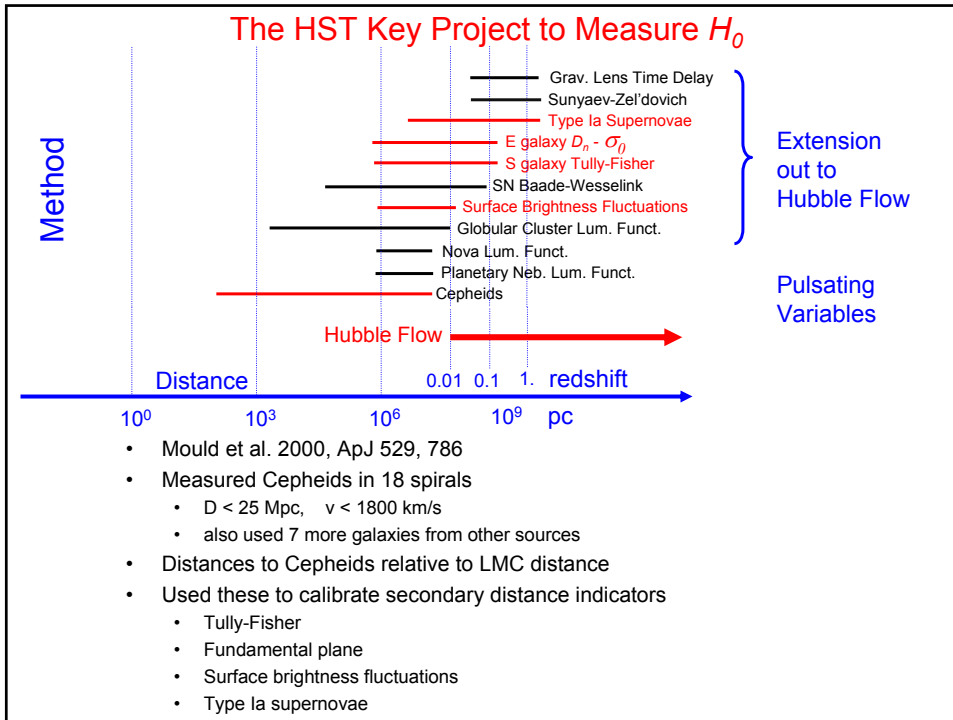
$$H_0 = 100h \text{ km s}^{-1} \text{ Mpc}^{-1}$$

- Hubble time

$$t_H = 1/H_0 = 9.78 \times 10^9 h^{-1} \text{ yr.}$$

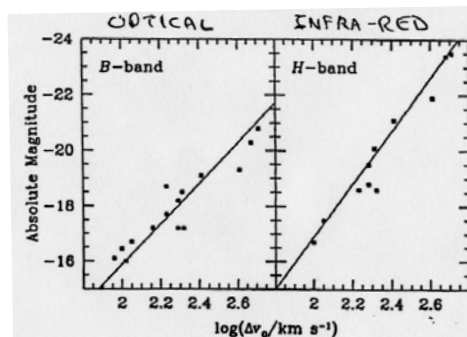
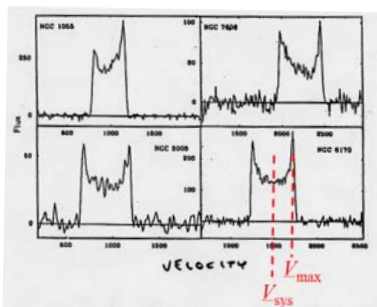
The HST Key Project to Measure H_0





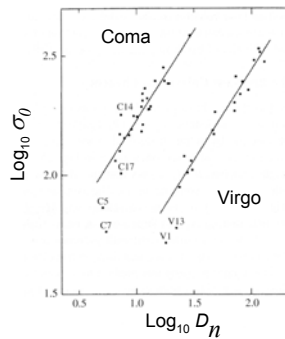
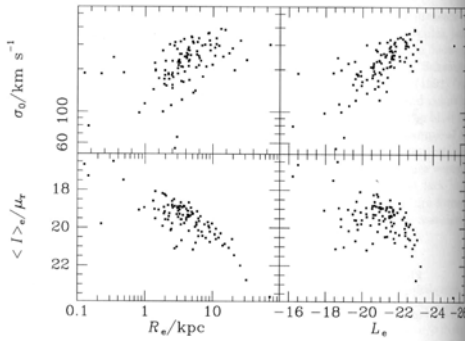
Tully-Fisher Relation

- L - v correlation
- for spiral galaxies, v easily measured using H I 21cm (radio) profiles.
- must apply $\sin i$ correction for inclination.
- infrared Tully-Fisher: IR measurements minimize scatter in L due to absorption \implies tighter correlation
- $F/L \rightarrow$ distance



E Galaxy Fundamental Plane The $D_n - \sigma_0$ relation

- Define:
 - D_n = angular diameter at which surface brightness reaches
 - $I_n = 20.75$ B-mag/arcsec²
- Observations show that linear size (in kpc) corresponding to D_n is tightly correlated with σ_0
- $D_n - \sigma_0$ relation combines radius, surface brightness and internal velocity dispersion
 - σ_0
 - The Fundamental Plane strikes again!
- Angular size = $D_n = (\text{linear size})/\text{distance}$
- 15% scatter in resulting distance to any one galaxy.



[CO Fig. 27.5]

Surface brightness fluctuations

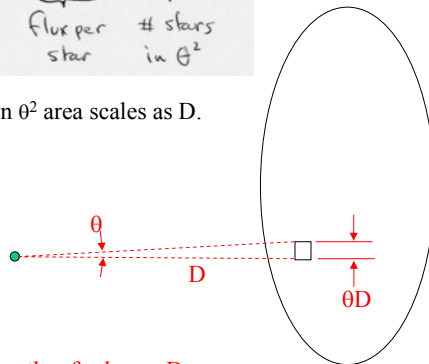
Same galaxy seen at any distance will have same surface brightness.

- N = number of stars in angular area θ^2
- n = number of stars per unit absolute area
- L = luminosity of each star
- D = distance

$$\text{Surface brightness} = I = \left(\frac{L}{4\pi D^2}\right) N = \underbrace{\left(\frac{L}{4\pi D^2}\right)}_{\text{flux per star}} \underbrace{n (D\theta)^2}_{\text{\# stars in } \theta^2} = \frac{Ln\theta^2}{4\pi}$$

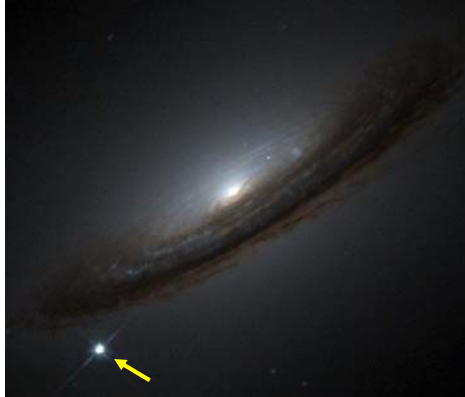
But standard deviation of number of stars in θ^2 area scales as D .

$$\begin{aligned} N &= n (D\theta)^2 \\ \sigma_N &= \sqrt{N} = \sqrt{n} D\theta \\ \sigma_I &= \left(\frac{L}{4\pi D^2}\right) \sigma_N = \frac{L n^{1/2} \theta}{4\pi D} \\ \frac{\sigma_I^2}{I^2} &= \frac{\langle L \rangle}{4\pi D^2} \end{aligned}$$



So surface brightness distributions look smoother for larger D .

Type Ia Supernovae



Core collapse supernovae

- Massive stars ($M > 8$ or $10 M_{sun}$)
- Wide range in $M \rightarrow$ wide range in L
- Not useful as “standard candles”

Type Ia supernovae

- White dwarf with $M \sim 1.4 M_{sun}$
- L can be precisely calibrated.
- Good standard candles.

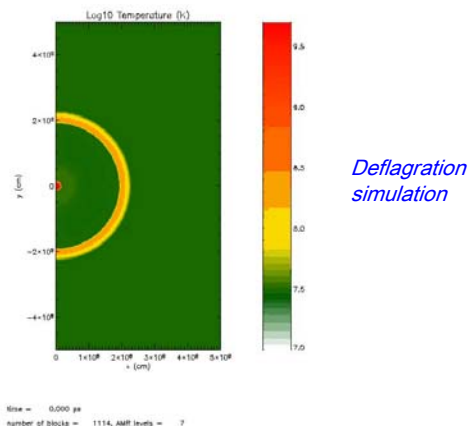
Type Ia Supernovae

- Something dumps too much mass onto white dwarf.
- Increased density \rightarrow runaway heating through C + C burning
- Heating rate faster than dynamical timescale
 - White dwarf cannot peacefully respond to pressure increase.
- *Deflagration*
 - leading to *detonation*?

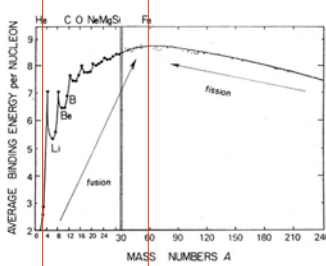
Type Ia Supernovae as “standard candles”.

- Always happens when mass goes just past limit for heating-cooling balance.
 - \rightarrow Supernova always has \sim same luminosity (factor 10).

- Get distance from
$$\text{Flux} = \frac{L}{4\pi r^2}$$



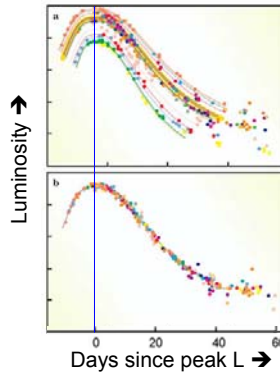
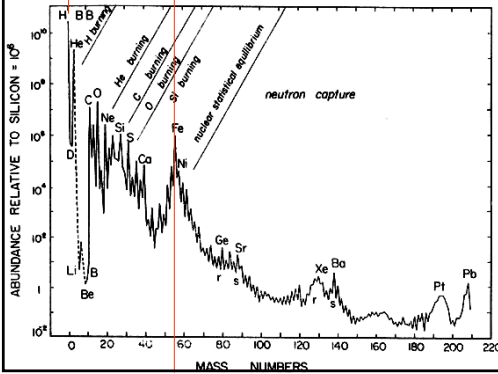
SN Ia as Standard Candles



Light output powered by radioactive decay:



Amount of Ni determines both luminosity *and* opacity.
 • So luminosity and fading timescale are correlated.

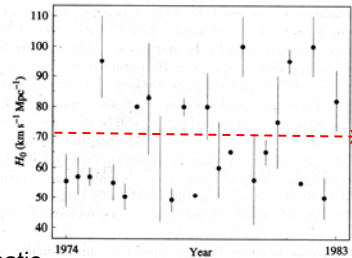


Observed range in *L* and fading timescale.

After correcting for *L* - timescale correlation.

The HST Key Project to Measure H_0

- Measured Distances to Cepheids.
 - relative to LMC distance.
- Used these to calibrate secondary distance indicators in same galaxies.

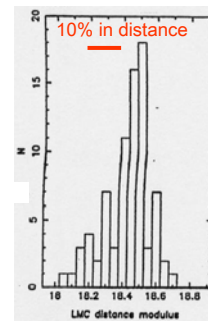


	value	random	systematic
Tully-Fisher	71	± 4	± 4
$D_n - \sigma_0$	78	8	10
Surface Brightness Fluct.	69	4	6
Type Ia SNe	68	2	2

Average: $H_0 = 71 \pm 6$ km/s/Mpc

- Uncertainties:
 - Correction for large scale flows
 - Distance to LMC.

Taken to be 50 kpc ± 6.5%

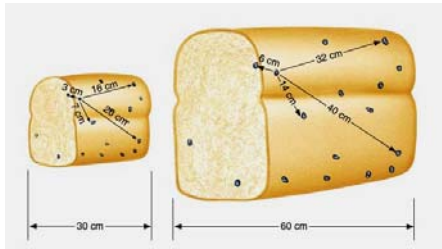


Distribution of published LMC distance moduli

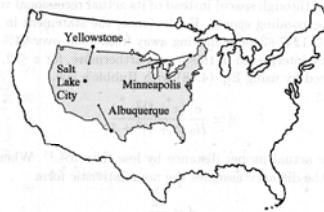
The Expanding Universe

Homogeneous * Isotropic

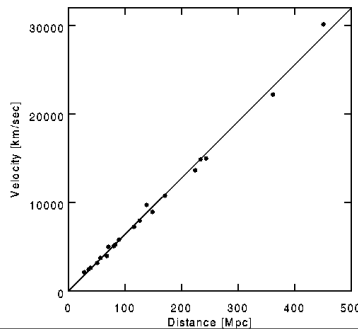
[CO 29.1]



As seen from a Wonder Bread raisin



As seen from Utah



The Cosmological Principle:

At any given time, the universe is the same everywhere.

Cosmological Principle: Universe is homogeneous & isotropic

Newtonian Cosmology

- Energy:

Kinetic + Potential = Total

$$\frac{1}{2}mv^2 - \frac{GM_r m}{r} = -\frac{1}{2}mkc^2\bar{\omega}^2 \quad [29.1]$$

$$\frac{1}{2}mv^2 - \frac{G \frac{4}{3}\pi r^3 \rho m}{r} = -\frac{1}{2}mkc^2\bar{\omega}^2$$

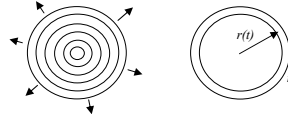
$$v^2 - \frac{8}{3}\pi G\rho r^2 = -kc^2\bar{\omega}^2 \quad [29.2]$$

$$\left(\left(\frac{1}{R} \frac{dR}{dt} \right)^2 - \frac{8}{3}\pi G\rho \right) R^2 = -kc^2 \quad [29.10]$$

Friedman eq'n:

- Nested, expanding shells

- Infinite series, all same density $\rho(t)$



- Follow single shell, mass m

[29.3]

- $r(t) = (\text{Scale factor}) \times (\text{co-moving coordinate})$

$$r(t) = R(t)\bar{\omega}$$

Define: Total Energy = $-\frac{1}{2}mkc^2\bar{\omega}^2$
Why???

Cosmological principle \rightarrow

For bound universe, each nested shell must simultaneously have KE $\rightarrow 0$

$$E = -G \frac{4}{3}\pi r^2 \rho m \propto m\bar{\omega}^2$$