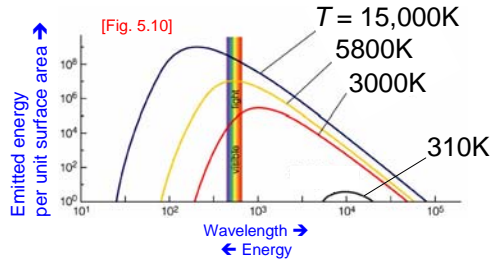


Basic Properties of Stars

Today's slides are on the web

- Composition
- Luminosity = L
- Mass
- Diameter = D



- Measure temperature T from spectrum.

- Total energy emitted *per unit surface area* = $E = \sigma T^4$

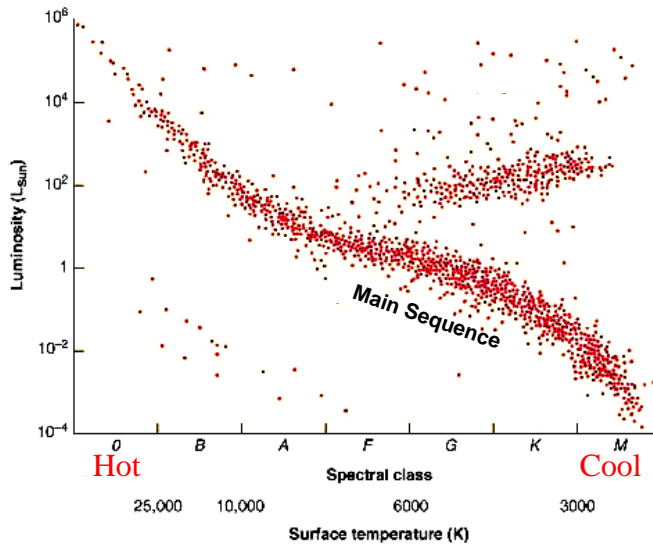
- Total energy from whole star:

$$L = E \times (\text{surface area}) = (\sigma T^4) \times (\pi D^2)$$

Measure L and T , solve for D

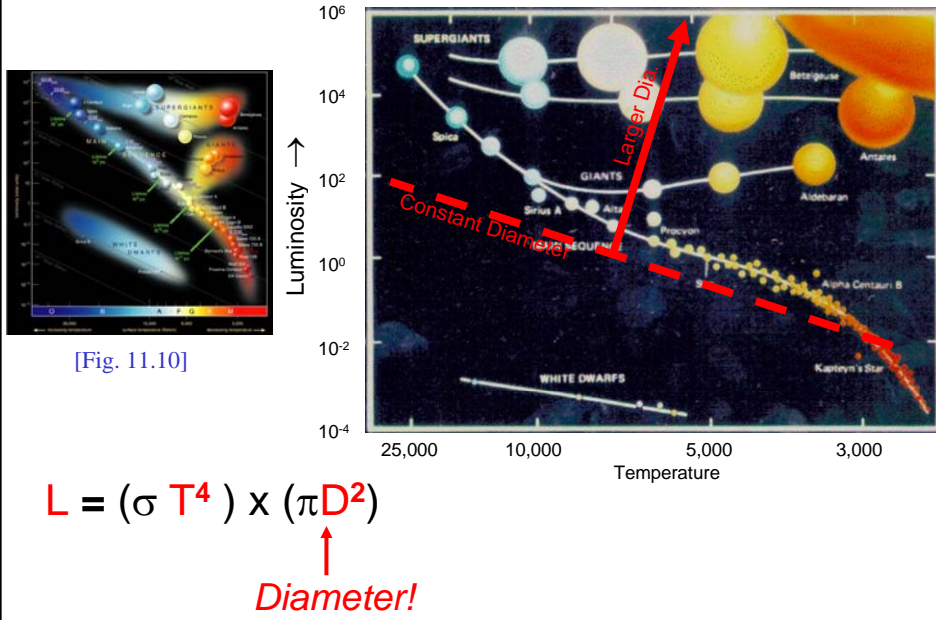
The Hertzsprung-Russell (H-R) Diagram [11.2]

- Representative sample of stars.
- Plot Luminosity vs. Surface Temperature.
- Most stars fall along “main sequence”.
- Any theory about how stars work has to explain this.

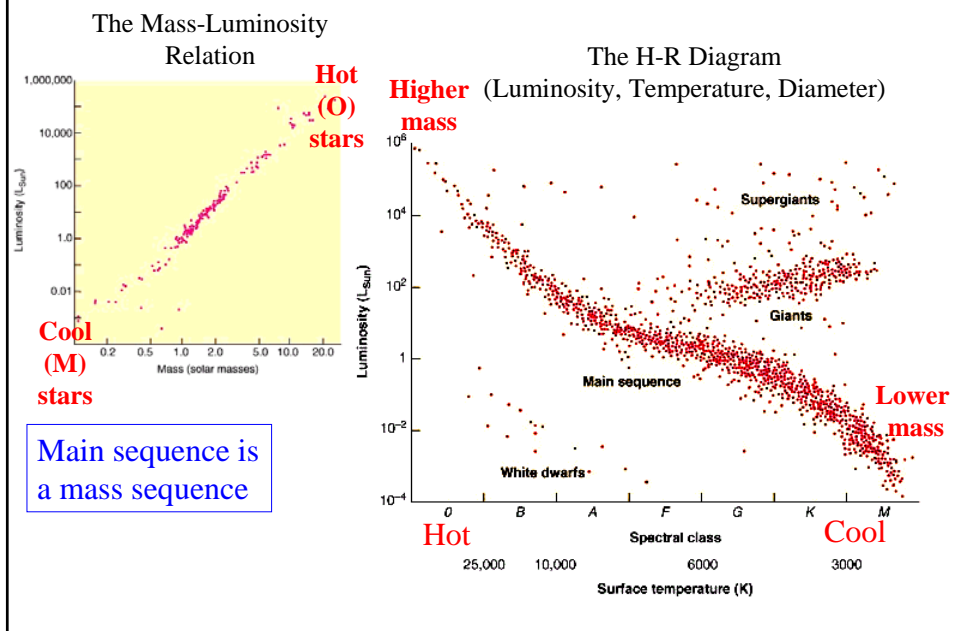


[see Fig. 11.10]

The H-R Diagram (with sizes)



Here's what we observe about stars.



What is inside the Sun?

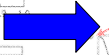
- Measure**
- Luminosity
 - Mass
 - Diameter
 - Chemical composition



- Infer**
(from our knowledge of Physics)
- Internal structure

What is inside other stars?

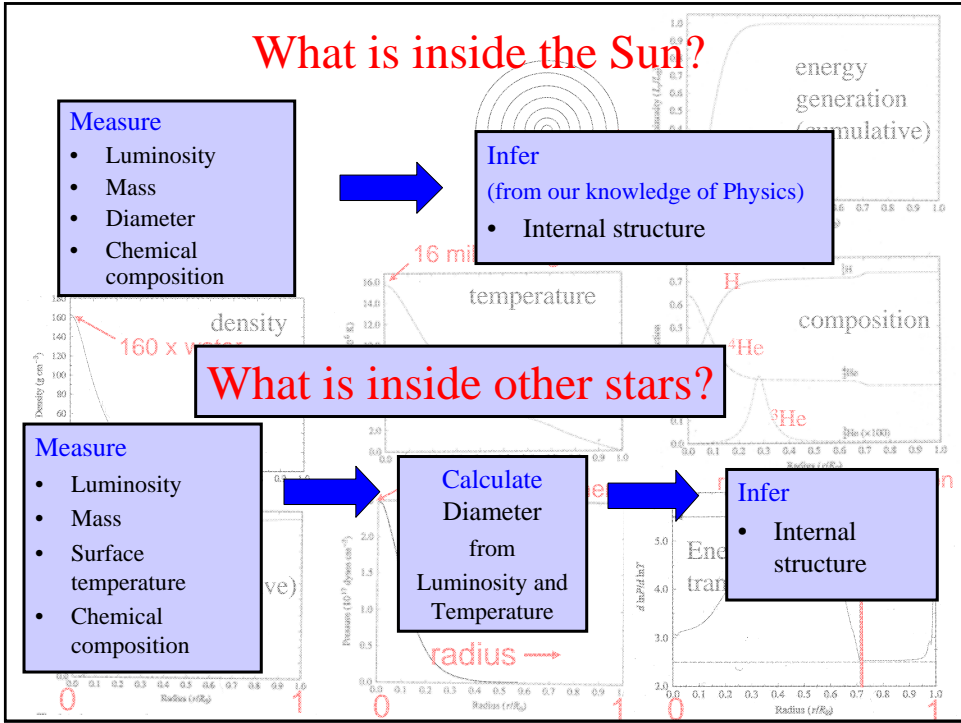
- Measure**
- Luminosity
 - Mass
 - Surface temperature
 - Chemical composition



- Calculate**
Diameter
from
Luminosity and
Temperature

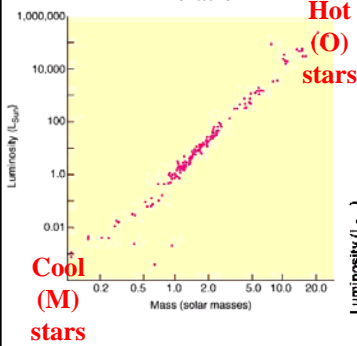


- Infer**
- Internal structure



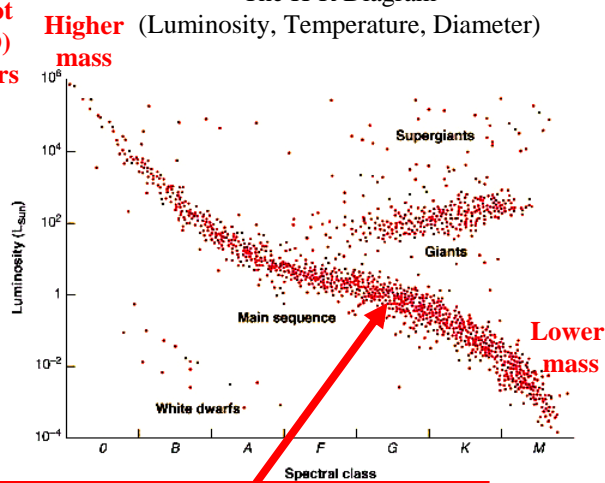
Here's what we observe about stars.

The Mass-Luminosity Relation

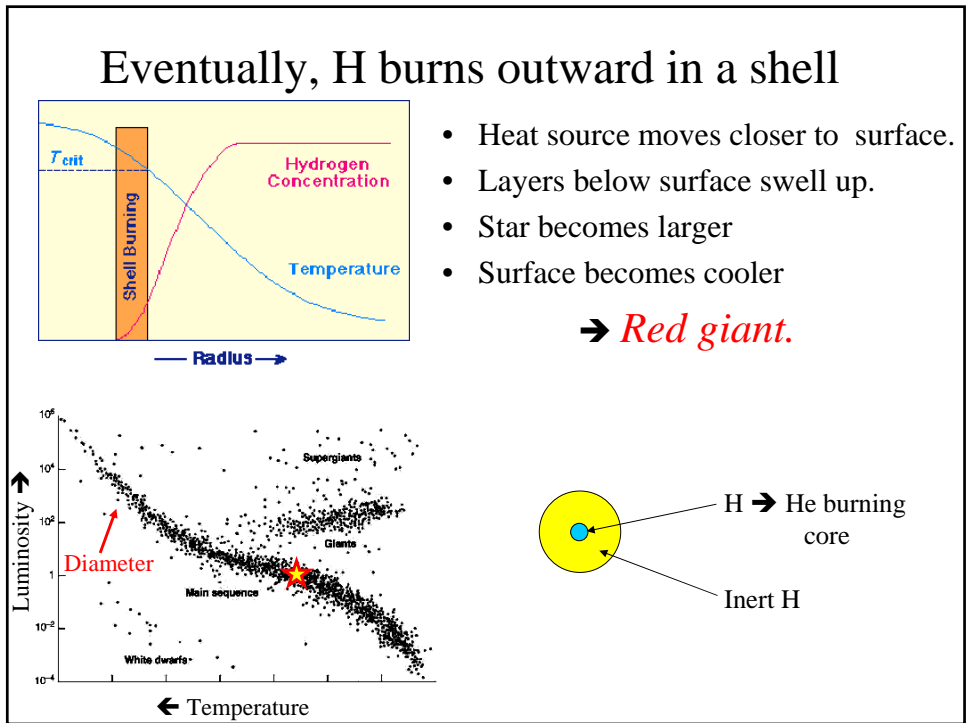
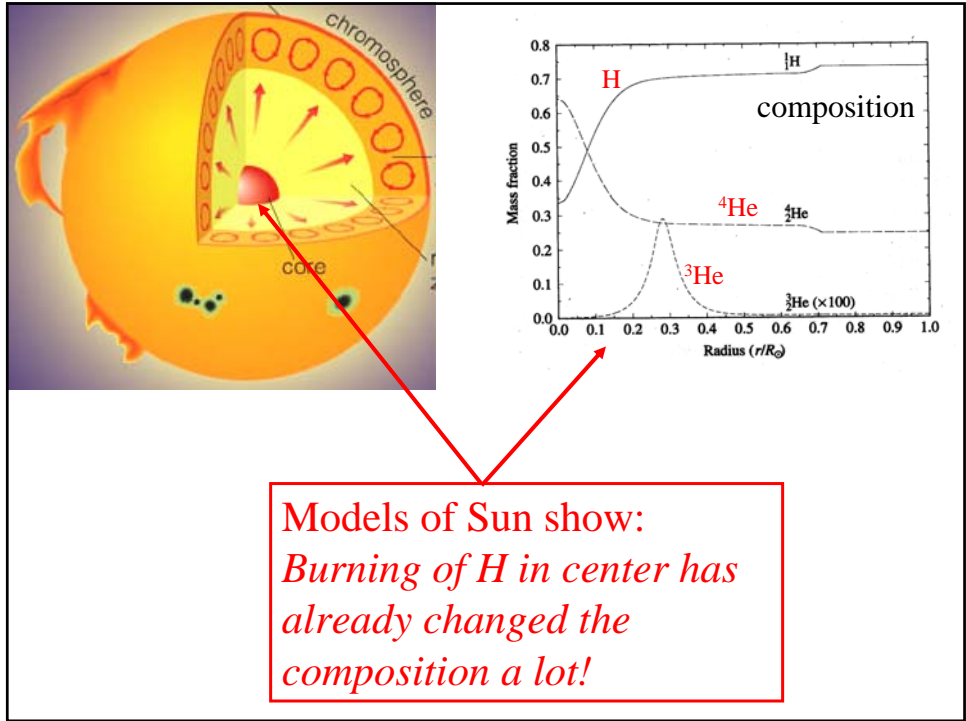


Main sequence is a mass sequence

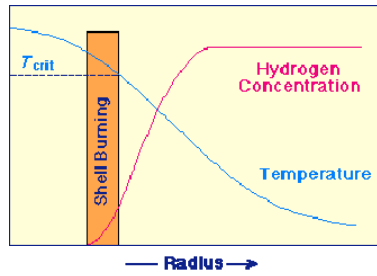
The H-R Diagram (Luminosity, Temperature, Diameter)



MAIN SEQUENCE:
Stars convert H into He in their cores.

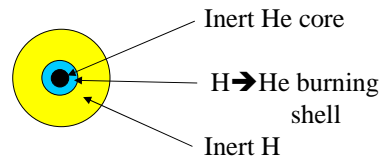
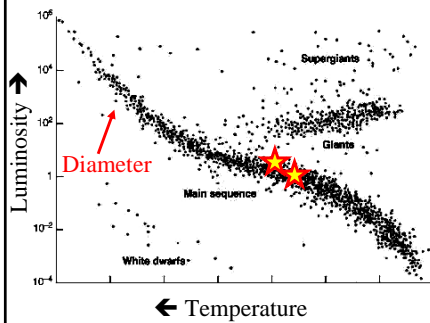


Eventually, H burns outward in a shell

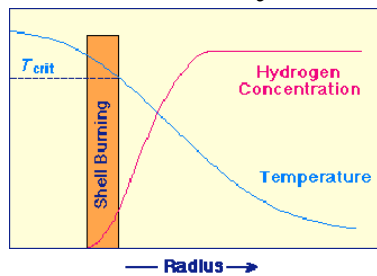


- Heat source moves closer to surface.
- Layers below surface swell up.
- Star becomes larger
- Surface becomes cooler

→ *Red giant.*

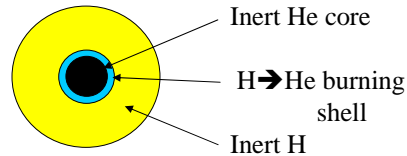
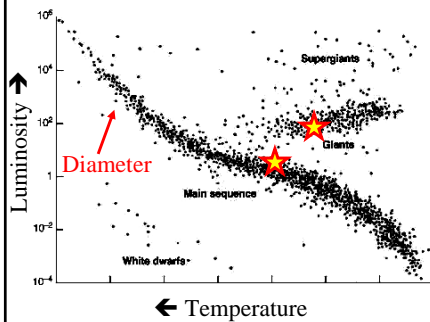


Eventually, H burns outward in a shell

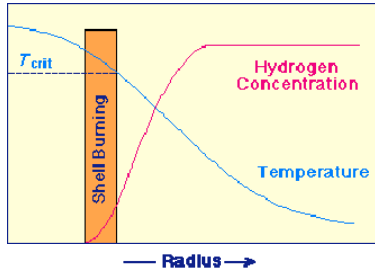


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