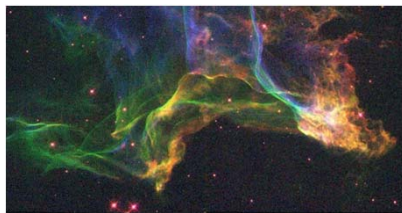


## Supernovae — Oct 21

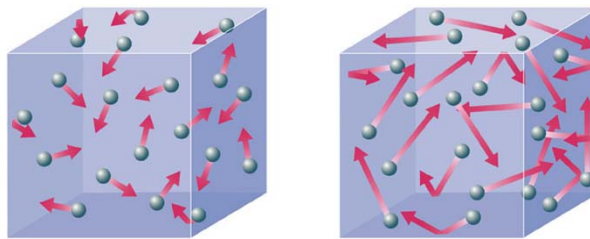


Cygnus Loop  
Supernova 20,000 yr ago

- Outline
  - Degenerate gas
  - Where were the elements in the baby made?
    - Carbon was made and expelled by giants
    - Iron was made in massive stars and expelled by supernovae
    - Heavier elements were made in supernovae & in giants, where there are free neutrons. Nuclei capture neutrons.
  - A massive star ( $M > 8M_{\text{sun}}$ ) ends as supernova
  - Elements produced in massive stars

## Normal/degeneracy pressure White dwarfs

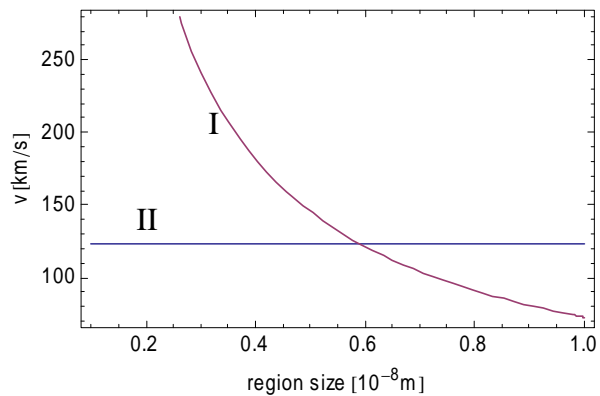
- Pressure on the walls of the box is caused by the gas hitting the walls. Gas transfers momentum to the walls.
- Mental picture: Marbles hit the walls; wall pushes back.



Longer arrows mean higher average speed.

## Pressure in a “degenerate gas”

- Normal gas
  - $PV = nkT$
  - Pressure is greater at higher temperature because the marbles are moving faster.  $mv^2=kT$
- Degenerate gas. If the gas is confined to a very small space, Newton’s 2<sup>nd</sup> law becomes invalid.
  - New laws of motion, called quantum mechanics, apply.
- Heisenberg’s uncertainty principle. Suppose a particle is allowed to move within a region of length  $x$ .
  - $mvx > h$
  - $\text{momentum } x > h$
  - $v$  is speed.  $h$  is Planck’s constant.
  - A particle must move if it is confined to a small space.
  - If you confine an electron to  $10^{-8}\text{m}$ , it moves at  $70\text{km/s}$ .




1. Plot shows the speed of a normal gas made of electrons with a temperature of  $1000\text{K}$  and a degenerate gas of electrons with a temperature of  $0\text{K}$ .
  - A. I is a NG. II is a DG.
  - B. I is a DG. II is a NG.


## Pressure in a “degenerate gas”

- Heisenberg’s uncertainty principle. Suppose a particle is allowed to move within a region of length  $x$ .  


$$\text{momentum } x > h$$
 A particle must move if it is confined to a small space.
- Pressure of a degenerate gas
  - $P V^{5/3} = \text{constant } n^{5/3}$
  - constant =  $h^2/m$
  - Pressure does not depend on temperature!!!
- End state of the sun
  - Without any energy production, the sun would normally shrink because gravity wins.
  - However, degeneracy pressure balances gravity. The sun is stable and cannot shrink.



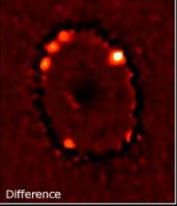
Large Magellanic Cloud



During Before



February 2, 2000




Difference

## Supernova 1987A

- Exploded in Large Magellanic Cloud
  - LMC is small galaxy that orbits our own Milky Way Galaxy.

Pre-existing circumstellar ring lit up first by photons from SN, now by blast wave from SN.



Feb '94 Sept '94 Mar '95 Feb '96

**Supernova 1987A**  
PRC97-03 - ST ScI OPO - January 14, 1997  
 J. Pui (NASA/GSFC), R. Kirshner (CIA) and NASA

## Guest star of 1054, Crab nebula



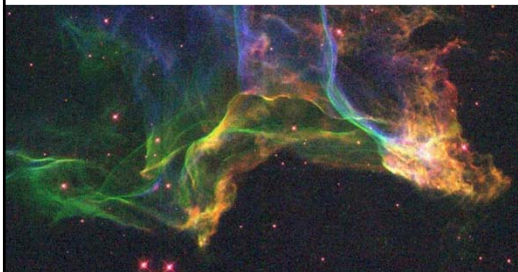
- Records of Sung Dynasty
  - In the first year of the period Chih-ho, ..., a guest star appeared several degrees SE of Thien-kuan. After more than a year it gradually became invisible.–p550.
- Gas expelled in 1054AD, still glowing, still moving.
- Other SN
  - 1572 Tycho
  - 1604 Kepler

## Supernova remnants

We expect one  
supernova in  
Milky Way every  
25-100 yrs.



Crab  
1,000 yrs old



Cygnus Loop  
20,000 yrs old.  
2500 LY away.



IC 443  
8000 yrs old

## Supernovae

- Explosion releases enormous energy
- Luminosity in photons temporarily exceeds that of whole galaxy full (100 billion) of stars.



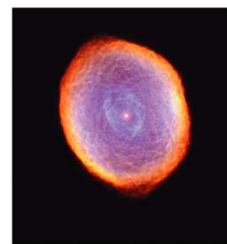
## What is a supernova? Why sun becomes a white dwarf, not a supernova

- In future double-shell burning sun, hot enough to burn  
 $3\text{}^4\text{He} \rightarrow \text{}^{12}\text{C}$
- When He exhausted, gravity wins, and core contracts.
- Electrons are so tight that they become degenerate.
- New source of pressure to resist gravity.
- Temperature not hot enough to burn carbon.

| Reaction   | Min. Temp.                  |
|--|-----------------------------|
| $4\text{}^1\text{H} \rightarrow \text{}^4\text{He}$                                      | $10^7\text{}^\circ\text{K}$ |
| $3\text{}^4\text{He} \rightarrow \text{}^{12}\text{C}$                                   | $2 \times 10^8$             |
| $\text{}^{12}\text{C} + \text{}^4\text{He} \rightarrow \text{}^{16}\text{O, Ne, Na, Mg}$ | $8 \times 10^8$             |
| $\text{Ne} \rightarrow \text{O, Mg}$   | $1.5 \times 10^9$           |
| $\text{O} \rightarrow \text{Mg, S}$  | $2 \times 10^9$             |
| $\text{Si} \rightarrow \text{Fe peak}$   | $3 \times 10^9$             |

## What is a supernova? Why sun becomes a white dwarf, not a supernova

- Sun has one more trick after He is exhausted in core.
  - Burn He in a shell
- Sun is not massive enough to shrink further and get hot enough to burn carbon.
 
$$T = M/R$$
  - Core is supported by pressure of degenerate electrons.
  - Temperature does not rise to burn anything else.
- End of the road: planetary nebula & white dwarf core



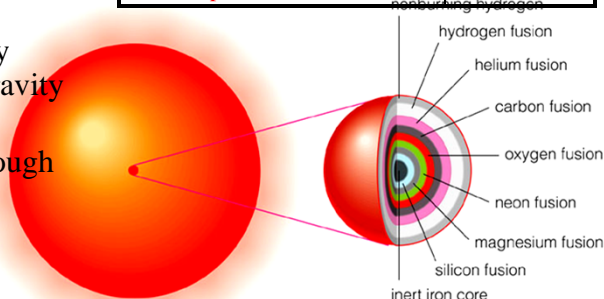
~~triple-alpha process~~

| Reaction  | Min. Temp. |
|---|------------|
| $4 \text{ } ^1\text{H} \rightarrow \text{ } ^4\text{He}$  | 10 MK      |
| $3 \text{ } ^4\text{He} \rightarrow \text{ } ^{12}\text{C}$   | 200 MK     |
| $\text{ } ^{12}\text{C} + \text{ } ^4\text{He} \rightarrow \text{ } ^{16}\text{O}, \text{Ne}, \text{Na}, \text{Mg}$ | 800 MK     |
| $\text{Ne} \rightarrow \text{O}, \text{Mg}$   | 1500MK     |
| $\text{O} \rightarrow \text{Mg}, \text{S}$  | 2000MK     |
| $\text{Si} \rightarrow \text{Fe peak}$  | 3000MK     |

## What is a supernova? Why massive star becomes a supernova

- History of a massive star
- During double-shell burning phase, hot enough to burn
 
$$3 \text{ } ^4\text{He} \rightarrow \text{ } ^{12}\text{C}$$
- When He exhausted, gravity wins, and core contracts.
- Temperature rises by larger amount b/c gravity is stronger.
- Temperature hot enough to burn carbon.
 
$$^4\text{He} + \text{ } ^{12}\text{C} \rightarrow \text{ } ^{16}\text{O}, \text{etc}$$

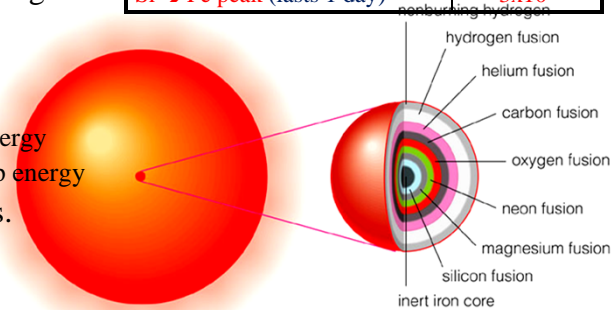
| Reaction  | Min. Temp.                     |
|---|--------------------------------|
| $4 \text{ } ^1\text{H} \rightarrow \text{ } ^4\text{He}$  | $10^7 \text{ } ^\circ\text{K}$ |
| $3 \text{ } ^4\text{He} \rightarrow \text{ } ^{12}\text{C}$   | $2 \times 10^8$                |
| $\text{ } ^{12}\text{C} + \text{ } ^4\text{He} \rightarrow \text{ } ^{16}\text{O}, \text{Ne}, \text{Na}, \text{Mg}$ | $8 \times 10^8$                |
| $\text{Ne} \rightarrow \text{O}, \text{Mg}$   | $1.5 \times 10^9$              |
| $\text{O} \rightarrow \text{Mg}, \text{S}$  | $2 \times 10^9$                |
| $\text{Si} \rightarrow \text{Fe peak}$  | $3 \times 10^9$                |



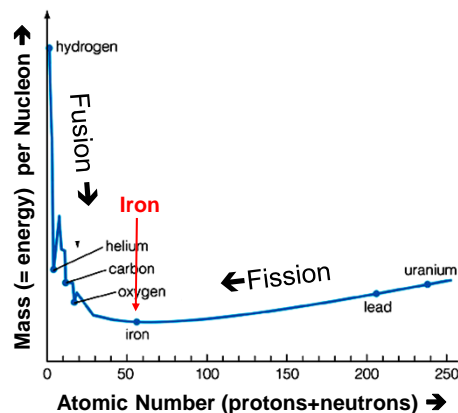
## What is a supernova? Why massive star becomes a supernova

- Hot enough to burn  $4\text{He} + {}^{12}\text{C} \rightarrow {}^{16}\text{O}$ , etc
- When C exhausted, gravity wins, and core contracts.
- Temperature rises.
- Temperature hot enough to burn neon.  ${}^{20}\text{Ne} + {}^4\text{He} \rightarrow {}^{24}\text{Mg}$
- Disaster with iron
  - Burning releases energy
  - Fusing iron takes up energy
- Gravity finally wins.

| Reaction   | Min. Temp.                   |
|--|------------------------------|
| $4\text{ }^1\text{H} \rightarrow {}^4\text{He}$  | $10^7\text{ }^\circ\text{K}$ |
| $3\text{ }^4\text{He} \rightarrow {}^{12}\text{C}$   | $2 \times 10^8$              |
| ${}^{12}\text{C} + {}^4\text{He} \rightarrow {}^{16}\text{O}, \text{Ne}, \text{Na}, \text{Mg}$ | $8 \times 10^8$              |
| $\text{Ne} \rightarrow \text{O}, \text{Mg}$  | $1.5 \times 10^9$            |
| $\text{O} \rightarrow \text{Mg}, \text{S}$   | $2 \times 10^9$              |
| $\text{Si} \rightarrow \text{Fe peak (lasts 1 day)}$   | $3 \times 10^9$              |



- Fusion of stable nuclei less massive than iron releases energy
- Fusion of stable nuclei more massive than iron requires input of energy and therefore cannot occur.



## What is a supernova? Why massive star becomes a supernova

- Disaster with iron
  - Burning releases energy
  - Fusing iron takes up energy
- Gravity finally wins.
- Star collapses in few seconds
- Rebounds as supernova
  - Reason for rebounding is topic of current research
- Expel outer layers.

| Reaction  | Min. Temp.             |
|---|------------------------|
| $4\ ^1\text{H} \rightarrow\ ^4\text{He}$  | $10^7\ ^\circ\text{K}$ |
| $3\ ^4\text{He} \rightarrow\ ^{12}\text{C}$   | $2 \times 10^8$        |
| $^{12}\text{C} +\ ^4\text{He} \rightarrow\ ^{16}\text{O},\ \text{Ne},\ \text{Na},\ \text{Mg}$ | $8 \times 10^8$        |
| $\text{Ne} \rightarrow\ \text{O},\ \text{Mg}$   | $1.5 \times 10^9$      |
| $\text{O} \rightarrow\ \text{Mg},\ \text{S}$  | $2 \times 10^9$        |
| $\text{Si} \rightarrow\ \text{Fe peak}$   | $3 \times 10^9$        |

